

Black Holes: Fermions at the Extremal Limit?

G. Ruppeiner*

Division of Natural Sciences

New College of Florida

Sarasota, Florida 34243-2109

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Abstract

I present exact results matching Kerr-Newman Black Hole thermodynamics in the extremal limit to the two-dimensional Fermi Gas. Two dimensions are consistent with the membrane paradigm of black holes. Key in the analysis is the thermodynamic Riemannian curvature scalar R , negative for most ordinary thermodynamic systems, including those near the critical point, but mostly positive for the Kerr-Newman Black Hole and the Fermi Gas.

The Kerr-Newman Black Hole (KNBH), with mass M , angular momentum J , and charge Q , has a well-established thermodynamic structure originated by Bekenstein [1] and Hawking [2]. Logically, this should be supported

*Electronic address: ruppeiner@ncf.edu

by some underlying microscopic picture, a "statistical mechanics" of black holes. But, since matter is expected to rapidly collapse to the central singularity, what entities could drive this statistics?

I argue that insight results from comparing known KNBH thermodynamics to standard statistical mechanical models. A major element here [3] is the thermodynamic Riemannian curvature scalar R , first evaluated for the KNBH by Åman, Bengtsson, and Pidokrajt [4]. In the sign convention of Weinberg [5], R is mostly positive for the KNBH, but mostly negative for ordinary thermodynamic systems, including most critical point models. An exception is the three-dimensional (3D) Fermi Gas, which Janyszek and Mrugała [6] found to have positive R . These authors also emphasized the general importance of the sign of R .

Here, I work out the 2D Fermi Gas and find some exact correspondences with the KNBH in the extremal limit, where the temperature goes to zero.

The KNBH entropy is [7]

$$S(M, J, Q) = \frac{1}{8} \left(2M^2 - Q^2 + 2\sqrt{M^4 - J^2 - M^2Q^2} \right), \quad (1)$$

in geometrized units [8], where S and J are in cm^2 , and M and Q are in cm .

The temperature T is given by

$$\frac{1}{T} \equiv \left(\frac{\partial S}{\partial M} \right)_{J,Q} = \frac{(K^2 + 2K + L^2) M}{4K}, \quad (2)$$

and the heat capacity

$$C_{J,Q} \equiv T \left(\frac{\partial S}{\partial T} \right)_{J,Q} = \frac{M^2 K (K^2 + L^2 + 2K)}{4(L^2 - 2K)}. \quad (3)$$

Here, the dimensionless simplifying variables

$$\{K, L\} \equiv \{\sqrt{1 - \alpha - \beta}, \sqrt{1 + \alpha}\}, \quad (4)$$

with $\alpha \equiv J^2/M^4$ and $\beta \equiv Q^2/M^2$.

To be in the *physical regime* with real, positive S and T requires

$$\alpha + \beta < 1. \tag{5}$$

Equality, $\alpha + \beta = 1$, has $K = T = 0$ and constitutes the *extremal limit*, unattainable by the third law of black hole thermodynamics [9]. In this limit, Eqs. (2) and (3) yield

$$C_{J,Q} = \frac{1}{16} M^3 L^2 T. \tag{6}$$

Imagine now the KNBH immersed in an *infinite, extensive* environment. The thermodynamic fluctuation probability is given by Einstein's formula

$$P \propto \exp(S_{tot}/k_B). \tag{7}$$

Here, S_{tot} is the total entropy of the universe and k_B is Boltzman's constant. If all three (M, J, Q) fluctuate, S_{tot} has no local maxima, and there are no stable states [10, 11]. But, if we formally restrict one of M , J , or Q as constant, reasoning it slow to fluctuate compared with the other two, stability is possible [11].

Consider (J, Q) fluctuating at constant M . This is stable for all thermodynamic states in the physical regime. I argued [11] that with M on the order of the Planck mass, such fluctuations might be physically relevant. A straightforward exercise [11, 12] results in the usual Gaussian approximation to Eq. (7):

$$P \propto \exp \left\{ -\frac{1}{2} \left[g_{22}(\Delta J)^2 + 2g_{23}\Delta J\Delta Q + g_{33}(\Delta Q)^2 \right] \right\}, \tag{8}$$

with

$$g_{\alpha\beta} \equiv - \left(\frac{8\pi}{L_p^2} \right) \frac{\partial^2 S}{\partial X^\alpha \partial X^\beta}, \quad (9)$$

$\{X^1, X^2, X^3\} \equiv \{M, J, Q\}$, ΔX^α the deviation of X^α from its value at maximum S_{tot} , the Planck length $L_p \equiv \sqrt{\hbar G/c^3}$, and \hbar , c , and G the usual physical constants. With an infinite, extensive environment, fluctuations depend only on the black hole thermodynamics. The environment merely sets the state about which fluctuations occur.

The constant multiplier in Eq. (9) converts S to S/k_B in real units [11], essential in Eq. (7). Discussion in the black hole literature tends to focus on where R is relatively large or small, and for this an overall multiplier for the metric is not particularly important. But, a fluctuation based metric whose R gets related to that in ordinary thermodynamic systems requires real units.

The quadratic form in Eq. (8):

$$(\Delta l)^2 \equiv g_{22}(\Delta J)^2 + 2g_{23}\Delta J\Delta Q + g_{33}(\Delta Q)^2, \quad (10)$$

is the line element for a Riemannian geometry of thermodynamics [3]. Its physical significance is clear from Eq. (8): *the less probable a fluctuation between two states, the further apart they are.*

Calculate R as follows [5]: the Christoffel symbols are

$$\Gamma_{\beta\gamma}^\alpha = \frac{1}{2}g^{\mu\alpha} (g_{\mu\beta,\gamma} + g_{\mu\gamma,\beta} - g_{\beta\gamma,\mu}), \quad (11)$$

with $g^{\alpha\beta}$ the inverse of the metric, and the comma notation indicating differentiation. The curvature tensor is

$$R_{\beta\gamma\delta}^\alpha = \Gamma_{\beta\gamma,\delta}^\alpha - \Gamma_{\beta\delta,\gamma}^\alpha + \Gamma_{\beta\gamma}^\mu \Gamma_{\mu\delta}^\alpha - \Gamma_{\beta\delta}^\mu \Gamma_{\mu\gamma}^\alpha, \quad (12)$$

and the Riemannian curvature scalar is

$$R = g^{\mu\nu} R_{\mu\xi\nu}^{\xi}. \quad (13)$$

R is independent of the choice of coordinate system, suggesting it is a fundamental measure of thermodynamic properties.

For an ordinary thermodynamic system, $|R|$ was interpreted [13] as proportional to the correlation volume ξ^d , where d is the system's spatial dimensionality and ξ its correlation length; see [3, 14] for review. A thermodynamic quantity, R , then reveals information normally thought to reside in the microscopic regime, ξ . Thus, R is interesting also in black hole physics, which has thermodynamic structures, but little microscopic information; see [15, 16, 17] for review.

By analogy, I interpret $|R|$ for black holes as the average number of correlated Planck areas on the event horizon [18]. This interpretation, putting the statistical degrees of freedom on the event horizon, fits the membrane paradigm for black holes [19].

For (J, Q) fluctuations, Eq. (13) yields

$$R = \frac{(K^5 + L^2 K^3 - 2K^3 - 2K^2 + 3L^2 K - 3K + 2)}{4\pi K (K^3 + L^2 K - K + 1)^2} \left(\frac{M_p}{M}\right)^2, \quad (14)$$

with Planck mass $M_p \equiv \sqrt{\hbar c/G}$. (In geometrized units, $L_p = M_p$.) Figure 1 shows R in the physical regime. It is real and positive, with a minimum of zero at $J = Q = 0$. R is regular except at the extremal limit, where, by Eqs. (2) and (14), its limiting form is

$$R = \frac{2M_p^2}{\pi M^3 L^2 T}. \quad (15)$$

The limiting product of curvature and heat capacity,

$$(R) \left(\frac{8\pi}{L_p^2} C_{J,Q} \right) = \left(\frac{2M_p^2}{\pi M^3 L^2 T} \right) \left(\frac{8\pi M^3 L^2 T}{16L_p^2} \right) = 1, \quad (16)$$

is a unitless, scale free constant independent of where we are on the extremal curve. The multiplier for $C_{J,Q}$ converts S in $C_{J,Q}$ to S/k_B in real units.

For comparison, Table 1 reviews thermodynamic curvature in ordinary systems. R is negative where attractive interactions dominate and positive where repulsive interactions dominate. Cases with weak interactions have $|R|$ small. There are three somewhat quirky cases, with both positive and negative R 's; however, these are not relevant here. (But, Cai and Cho [32] connected phase transitions in BTZ black holes to the "quirk" in R for the Takahashi gas.) The only known simple situation with positive R diverging at low temperature is the Fermi gas, and I will turn to it for insight below.

In ordinary thermodynamic systems, it is traditional to pull a constant volume V out of the line element Eq. (10), giving R units of volume. ("Volume" depends on dimension; for 2D systems, it is area.) But, the KNBH has no fixed scale to pull out, and so for it we work with the full dimensionless line element Eq. (10). Its R is dimensionless, and is interpreted [18] as the number, rather than the volume, of correlated volumes. For ordinary thermodynamic systems, whether or not the fixed V is left in the line element makes no physical difference.

For the cold 3D Fermi Gas, R seems to diverge [6] as $T^{-3/2}$, and not as T^{-1} in Eq. (15) for the KNBH. This motivates me to work out the 2D Fermi Gas. By the reasoning leading to Eq. (8.1.3) of [12], the 2D Fermi gas has

thermodynamic potential

$$\phi(1/T, -\mu/T) = p/T = k_B g \lambda^{-2} f_2(\eta), \quad (17)$$

with pressure p , $\eta \equiv \exp(\mu/k_B T)$, chemical potential μ , thermal wavelength $\lambda \equiv h/\sqrt{2\pi m k_B T}$, particle mass m , weight factor $g \equiv (2s + 1)$, particle spin s , and

$$f_l(\eta) \equiv \frac{1}{\Gamma(l)} \int_0^\infty \frac{x^{l-1} dx}{\eta^{-1} e^x + 1}. \quad (18)$$

I use obvious fluid units for all quantities, including S and T . The integral in Eq. (18) converges for $f_2(\eta)$, and yields $f_1(\eta) = \ln(1 + \eta)$. $f_0(\eta)$ and $f_{-1}(\eta)$ follow from $f_1(\eta)$ using the recurrence relation $f_{l-1}(\eta) = \eta f'_l(\eta)$ [12].

Define the heat capacity at constant particle number N and constant area A by

$$C_{N,A} \equiv T \left(\frac{\partial S}{\partial T} \right)_{N,A} = N k_B [2f_2(\eta)/f_1(\eta) - f_1(\eta)/f_0(\eta)]. \quad (19)$$

The second equality is by Problem 8.10.ii of [12]. The methods of [12] now yield the limiting low T expression

$$\frac{C_{N,A}}{A k_B} = \frac{2\pi^3 g m k_B T}{3h^2}. \quad (20)$$

Evaluating R with Eq. (6.31) of [3] yields

$$R = -g^{-1} \lambda^2 \left\{ \frac{-2f_2(\eta)f_0(\eta)^2 + f_1(\eta)^2 f_0(\eta) + f_{-1}(\eta)f_1(\eta)f_2(\eta)}{[f_1(\eta)^2 - 2f_0(\eta)f_2(\eta)]^2} \right\}. \quad (21)$$

Numerical evaluation over the physical range $-\infty < \mu < +\infty$ and $0 < T < \infty$ indicates R is always positive. The methods of [12] yield the limiting low T expression:

$$R = \frac{3h^2}{2\pi^3 g m k_B T}. \quad (22)$$

The limiting T dependences of $C_{N,A}$ and R match the corresponding KNBH quantities Eqs. (6) and (15). This connection to a 2D model is consistent with the membrane paradigm of black holes [19]. Furthermore, the limiting product of curvature and heat capacity,

$$\left(\frac{R}{A}\right) \left(\frac{C_{N,A}}{k_B}\right) = \left(\frac{3h^2}{2\pi^3 g m k_B T A}\right) \left(\frac{2\pi^3 g m k_B T A}{3h^2}\right) = 1, \quad (23)$$

is a unitless, scale free constant independent of density. The factor A below R undoes the traditional pulling out of A in the ordinary thermodynamic line element. R/A here is analogous to R for the KNBH. The constant products Eqs. (16) and (23) are equal, remarkable for systems apparently so different.

Note a key difference. The KNBH entropy Eq. (1) does not go to zero in the extremal limit, as it does for the 2D Fermi Gas with its unique ground state. Resolution probably requires a more sophisticated Fermi gas model.

Results above are for (J, Q) fluctuations. Repeating the exercise for the quite different (M, Q) and (M, J) fluctuations [18], leads, remarkably, to the same limiting form for R as Eq. (15), and the same match to the 2D Fermi gas.

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System	R sign	Divergence
1D Ising ferromagnet [20, 21]	–	$T \rightarrow 0$
critical region [3, 13, 22]	–	critical point
mean-field theory [21]	–	critical point
van der Waals [3, 22]	–	critical point
Ising on Bethe lattice [23]	–	critical point
Ising on 2D random graph [14, 24]	–	critical point
spherical model [14, 25]	–	critical point
3D Bose gas [6]	–	$T \rightarrow 0$
self-gravitating gas [26]	–	unclear
1D Ising antiferromagnet [20, 21]	–	$ R $ small
Tonks gas [27]	–	$ R $ small
pure ideal gas [13]	0	$ R $ small
ideal paramagnet [20, 21]	0	$ R $ small
multicomponent ideal gas [28]	+	$ R $ small
Takahashi gas [27]	+/-	$T \rightarrow 0$
finite 1D Ising ferromagnet [29]	+/-	$T \rightarrow 0$
1D Potts model [14, 30]	+/-	critical point
3D Fermi gas [6]	+	$T \rightarrow 0$
3D Fermi paramagnet [31]	+	$T \rightarrow 0$

Table 1. Signs of R and where it diverges for ordinary thermodynamic systems. All signs are put into the sign convention of Weinberg [5]. A designation ” $|R|$ small” means a value on the order of the volume of an intermolecular spacing or less.

FIGURE CAPTION

Figure 1. $R(M/M_p)^2$ as a function of J/M^2 and Q/M for (J, Q) fluctuations. R is real, positive, and regular in the physical regime, and diverges as T^{-1} at the extremal limit.

Figure 1

