

Comments on *Unified dark energy and dark matter from a scalar field different from quintessence*

Luis P. Chimento^{1, a} and Mónica Forte^{1, b}

¹ *Departamento de Física, Facultad de Ciencias Exactas y Naturales,
Universidad de Buenos Aires, Ciudad Universitaria, Pabellón I, 1428 Buenos Aires, Argentina*
(Dated: May 23, 2018)

In a recent paper by C. Gao, M. Kunz, A. Liddle and D. Parkinson [1] the unification of dark matter and dark energy was explored within a theory containing a scalar field of non-Lagrangian type. This scalar field, different from the classic quintessence, can be obtained from the scalar field representation of an interacting two-fluid mixture described in the paper by L.P. Chimento and M. Forte [2].

I. INTRODUCTION

Due to the additivity of the stress-energy tensor is possible to describe the quintessence field in terms of two interacting fluids, namely, stiff matter and vacuum energy. Besides, the total energy-momentum conservation is equivalent to the Klein-Gordon equation [3]- [4]. This description was generalized assuming a mixture of two fluids with constant equations of state, which interact between them in a flat Friedmann-Robertson-Walker model [2]. There, a scalar field ϕ representation of the mixture of two interacting fluids was introduced by imposing the condition that $\dot{\phi}^2$ is an appropriate linear combination of the energy densities ρ_1 and ρ_2 of both fluids. Under these conditions, the resulting model was called exotic quintessence.

II. EXOTIC QUINTESSENCE

We present the exotic quintessence model developed in [2] including two extra matter components, with energy densities ρ_3 and ρ_4 , which separately satisfy their own equations of conservation. In this case, the Einstein equations read:

$$3H^2 = \rho_1 + \rho_2 + \rho_3 + \rho_4, \quad (1)$$

$$\dot{\rho}_1 + \dot{\rho}_2 + 3H[(1 + w_1)\rho_1 + (1 + w_2)\rho_2] = 0, \quad (2)$$

$$\dot{\rho}_3 + 3H(1 + w_3)\rho_3 = 0, \quad (3)$$

$$\dot{\rho}_4 + 3H(1 + w_4)\rho_4 = 0, \quad (4)$$

where we have adopted constant equations of state $w_n = p_n/\rho_n$ for each fluid with $n = 1,2,3,4$. Units are chosen so that the gravitational constant is set to $8\pi G = 1$ and $c = 1$. We introduce a scalar field ϕ representation of the interacting two first fluids associating $\dot{\phi}^2$ with the

following linear combination of ρ_1 and ρ_2

$$\dot{\phi}^2 = (1 + w_1)\rho_1 + (1 + w_2)\rho_2. \quad (5)$$

From Eqs. (2) and (5) we obtain the total energy density and pressure of the four-fluid mixture, and the dynamical equation for the scalar field

$$\rho = \frac{\dot{\phi}^2}{1 + w_1} + \frac{w_1 - w_2}{1 + w_1} \rho_2 + \rho_3 + \rho_4, \quad (6)$$

$$p = \frac{w_1 \dot{\phi}^2}{1 + w_1} - \frac{w_1 - w_2}{1 + w_1} \rho_2 + w_3 \rho_3 + w_4 \rho_4, \quad (7)$$

$$\ddot{\phi} + \frac{3}{2}(1 + w_1)H\dot{\phi} + \frac{w_1 - w_2}{2} \frac{\dot{\rho}_2}{\dot{\phi}} = 0. \quad (8)$$

These equations define the exotic quintessence model.

III. CONCLUSIONS

Choosing $w_1 = 0$, $w_3 = 0$, $w_4 = 1/3$, and making the substitutions $\phi \rightarrow \phi/\sqrt{2}$ and $-w_2\rho_2 \rightarrow \Lambda(\phi)$, the Einstein and the exotic quintessence equations are reduced to the expressions (21)-(23) of the paper [1]. On the other hand, if we impose the integrability condition mentioned in [2]

$$\dot{\rho}_2 + A\dot{\phi}\rho_2 = 0, \quad (9)$$

on the field equation (8) with a constant A , we get $\rho_2 \propto \Lambda(\phi) \propto \exp(-A\phi)$ after integrate the Eq. (9). Then, the exotic quintessence model with the integrability condition (9) leads to an exotic scalar field driven by an exponential potential and thus to obtain the model investigated in Ref. [1]. As was already pointed in Ref. [2], this model is integrable and can be solved exactly.

[1] C. Gao, M. Kunz, A. R. Liddle and D. Parkinson, Phys. Rev. D **81**, 043520 (2010) [arXiv:0912.0949 [astro-ph.CO]].
[2] L. Chimento and M. Forte, Phys. Lett. B **666**, 205 (2008) [arXiv:0706.4142 [astro-ph]].

[3] L. P. Chimento, Phys. Rev. D **65**, 063517 (2002).
[4] L. P. Chimento, Phys. Lett. B **633**, 9 (2006) [arXiv:gr-qc/0503049].