

Radio Spectra of a Sample of X-ray Selected BL Lacs

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Abstract We present simultaneous multifrequency radio observations for a complete subsample of 26 XBLs from the *Einstein* Extended Medium-Sensitivity Survey, obtained with the Very Large Array (VLA). Spectra are computed using fluxes at 20, 6 and 3.6 cm. Unlike many radio selected samples, the EMSS did not impose any criterion on the radio spectrum to identify BL Lac objects. It is therefore possible to investigate the intrinsic radio spectral slope distribution and to determine the effect produced by this selection criterion. We find that 15% of the observed objects do not meet the flat-spectrum criterion imposed on some other BL Lac samples. A dataset that includes non-simultaneous data (that are also taken with different VLA configurations) shows an even higher percentage of steep spectrum sources. This effect can be ascribed to a larger fraction of extended flux detected with the more compact VLA configuration. Possible biases introduced by the flat-radio-spectrum criterion in the radio-selected BL Lac samples cannot explain the discrepancies observed in the evolutionary properties of Radio and X-ray selected samples of BL Lacs.

Key words. BL Lacertae objects: general; Radio continuum: galaxies

1. Introduction

BL Lac objects are an enigmatic class of active galactic nuclei (AGN) characterized by strong radio, optical and X-ray variability, relatively high optical and radio polarization and featureless optical spectra (e.g. Urry & Padovani, 1995). The properties of the class can be explained in terms of the *relativistic beaming* scenario wherein the observed emission is dominated by Doppler-boosted non-thermal radiation from a relativistic jet aligned with the line of sight (Blandford & Rees, 1978, Antonucci & Ulvestad, 1985).

Owing to their featureless spectra and lack of UV excess in many BL Lacs, optical search techniques (e.g. by excess colors, emission line strength), used to search for other AGN, failed to find BL Lacs in large number (e.g. Fleming et al., 1993). Since BL Lacs are both radio-loud and X-ray-loud, surveys in these frequency bands discover them with high efficiency (e.g. the “1 Jansky” sample: Stickel et al., 1991; the *Einstein* Extended Medium-Sensitivity Survey (EMSS) sample: Morris et al., 1991; Rector et al., 2000; the Deep X-ray Radio Blazar Survey (DXRBS) sample: Perlman et al., 1998, Landt et al. 2001; the ROSAT All Sky Survey (RASS) sample: Bade et al., 1998; the “sedentary” survey: Giommi, Menna &

Padovani, 1999; the Radio Emitting X-ray (REX) survey sample: Caccianiga et al., 1999, Caccianiga et al., 2002). This splits BL Lacs into two empirical subclasses, namely *radio-selected BL Lacs* (RBLs) and *X-ray-selected BL Lacs* (XBLs). Padovani & Giommi (1995) proposed to distinguish these two classes on a more physical basis, according to the spectral energy distribution (SED). For ease of description here we will refer to RBL and XBL only.

RBLs and XBLs show different behaviors with respect to properties like cosmological evolution, polarization, variability, core-to-extended radio flux ratio, extended radio morphology, spectral features (e.g. Wolter et al., 1994; Stickel et al., 1991; Morris et al., 1991; Jannuzi, Smith & Elston, 1993, 1994; Perlman & Stocke, 1993 to mention just a few) which have not been satisfactorily explained yet. This led to the suggestion that the observed discrepancies are at least partially caused by selection effects due to different criteria used in radio and X-ray surveys.

In particular, the standard selection technique for RBLs requires a flat-radio-spectrum criterion for which the radio spectral slope must be less than 0.5 ($\alpha_r \leq 0.5$; $S_\nu \propto \nu^{-\alpha_r}$). This criterion has not been used for classical X-ray selected samples of BL Lacs like the EMSS. If surveys with this criterion select only part of the BL Lac population, the RBL samples could be biased and the dif-

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ferences between RBLs and XBLs like the ones mentioned above could be at least partially explained.

It is already known that a number of objects with steep spectra ($\alpha_r > 0.5$), classified as radio galaxies in radio surveys like the “1 Jansky” (Owen, Ledlow, & Keel, 1996, Perlman et al. 1996), otherwise meet the BL Lacs selection criteria, based on optical properties (equivalent width of emission lines; Ca II break contrast) of Marchã et al. (1996), and have broadband properties (radio luminosity and overall spectral energy distributions) that agree with those of BL Lacs (e.g. Rector et al., 2000).

To test if the flat–radio–spectrum criterion really introduces biases in the BL Lacs selection and to estimate the incompleteness degree of RBLs samples, we studied the radio spectral indices of a complete subsample of XBLs in which the identification criteria were not based on the radio spectral properties. Previous multi-frequency studies had already been performed on XBLs. For instance, Laurent-Muehleisen et al. (1993) collected non-simultaneous data for several objects of this class, that however do not make a complete sample. We prefer to deal with the largest subsample for which we could have both simultaneous and non-simultaneous data.

The sources of this subsample are extracted from the EMSS (Gioia et al., 1990; Stocke et al., 1991), a catalog whose BL Lacs sample has been thoroughly studied for completeness (Rector, Stocke & Perlman, 1999). A pilot study on the radio spectra of 8 XBLs from the EMSS and the *HEAO-1 A-2* all-sky survey (Piccinotti et al., 1982) was performed by Stocke et al. (1985). Three objects showed a spectral slope which exceeds the limit for flat spectrum and two more were marginally steep. Preliminary analysis of the EMSS BL Lacs sample, by using non-simultaneous data at 6 and 20 cm from the EMSS and the *NRAO VLA Sky Survey* (NVSS: Condon et al., 1998) respectively, showed that about 30% of the objects have a spectral slope steeper than 0.5. This is the motivation that led us to investigate the EMSS sources by obtaining simultaneous spectra, which are not affected by variability.

We briefly describe the sample of objects used for this work in §2; in §3 we present the data reduction process as well as flux densities and radio spectra for each object; in §4 we show the results obtained; our conclusions are summarized in §5. Throughout the paper we used $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $q_0 = 0$, but no conclusions are dependent upon this choice of cosmology.

2. The Sample

We restrict the analysis to the EMSS north of $\delta = -20^\circ$ in order to ensure 100% identification rate. The revised EMSS BL Lac sample (Rector et al., 2000) then contains 36 BL Lacs from which we observed a subsample of 22 objects, named hereafter CWSR. The sources were selected imposing $f_{6\text{cm}} \geq 1 \text{ mJy}$ in the EMSS dataset and celestial positions that allowed observations at the Very

Name	Obs. date	z
MS0122.1+0903	16 Jun 2000	0.339*
MS0158.5+0019	16 Jun 2000	0.299*
MS0205.7+3509	16 Jun 2000	> 0.351 ⁺
MS0257.9+3429	16 Jun 2000	0.245*
MS0317.0+1834	Jun 1984	0.190 [†]
MS0419.3+1943	16 Jun 2000	0.512*
MS0607.9+7108	16 Jun 2000	0.267*
MS0737.9+7441	16 Jun 2000	0.315*
MS0922.9+7459	13 Jun 2000	0.638*
MS0950.9+4929	13 Jun 2000	...
MS1019.0+5139	13 Jun 2000	0.141*
MS1050.7+4946	13 Jun 2000	0.140*
MS1207.9+3945	Jun 1984	0.616 [†]
MS1221.8+2452	13 Jun 2000	0.218*
MS1229.2+6430	13 Jun 2000	0.164*
MS1235.4+6315	Jun 1984	0.297 [†]
MS1402.3+0416	Jun 1984	0.344: [†]
MS1407.9+5954	13 Jun 2000	0.495*
MS1443.5+6348	13 Jun 2000	0.299*
MS1458.8+2249	13 Jun 2000	0.235*
MS1534.2+0148	13 Jun 2000	0.312*
MS1552.1+2020	13 Jun 2000	0.222*
MS1757.7+7034	13 Jun 2000	0.407*
MS2143.4+0704	16 Jun 2000	0.237*
MS2336.5+0517	16 Jun 2000	0.74:. [†]
MS2347.4+1924	16 Jun 2000	0.515*

(*) Redshift from Morris et al. (1991)

(⁺) Redshift from Watson et al. (2004)

([†]) Redshift from Rector et al. (2000)

Table 1. The CWSR sample.

Large Array (VLA¹) during June 2000. The resulting subsample is therefore not biased with respect to radio spectral index.

The VLA observations were performed on June 13 and June 16 in the L ($\lambda = 20 \text{ cm}$), C ($\lambda = 6 \text{ cm}$) and X ($\lambda = 3.6 \text{ cm}$) frequency bands. The observations were carried out in the C configuration. Thus, the maps have an angular resolution equal to $16''.3$, $5''.1$ and $3''$ at 20, 6 and 3.6 cm respectively. Each object was observed for ~ 10 minutes per frequency, obtaining a sensitivity of about $0.05 \text{ mJy beam}^{-1}$. Flux densities were bootstrapped from 3C 286 (June 13) and 3C 48 (June 16) (Baars et al. 1977)

If we include in the CWSR sample the 4 EMSS objects published in Stocke et al. (1985), observed simultaneously on June 1984 at all the available wavelengths (20, 18, 6, 2.0 and 1.3 cm) with the VLA in the C/D hybrid configuration, we have 26 sources in total. These sources are listed in Table 1: the first column indicates the source names, the second one the observation dates and the third one the redshift values.

¹ The VLA is a facility of the National Radio Astronomy Observatory. The National Radio Astronomy Observatory (NRAO) is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

For the same 26 objects we also have non-simultaneous data at 20 and 6 cm, respectively from NVSS and EMSS archives.

3. Data analysis

3.1. Radio maps

Radio maps for the 22 CWSR sources observed on June 2000 were produced in *ATPS* with a standard procedure that includes: removal of RFI from (u,v) data caused by incidental factors during the observations, calibration of all sources in amplitude and in phase, creation of a “cleaned” map of brightness distribution by using a CLEAN algorithm. Particular care was taken at 20cm to look and correct for contaminating sources.

For a more stable phase calibration, a calibrator source was observed before and after each source observation. Whenever possible, self-calibration was also applied.

We adopted a uniform weight function. We did some comparison with maps derived from natural weighting function and no significant differences were observed.

The resulting maps do not show morphological structures that are worth publishing. All sources are fitted by Gaussians (even if sometimes with widths larger than the beam size).

3.2. Flux densities

3.2.1. Simultaneous flux densities

From the “cleaned” maps we computed flux densities and their associated rms errors by assuming the source signal as a Gaussian. We measured both the peak flux density (F_{peak}) and the total flux density (F_{tot}). We assume that F_{peak} represents the emission from the source core and we define the extended flux F_{ext} as the difference between F_{tot} and F_{peak} .

In Table 2 F_{tot} , F_{peak} and their associated rms errors are listed for the CWSR sample. By comparing the peak and total flux densities, we confirm that this sample and the C data we obtained follow the BL Lac trend to be compact: we found that the maximum value of the ratio F_{tot}/F_{peak} is 1.98 at 20 cm. In §4.1 we better analyze this aspect.

The noise for each map was determined from the rms fluctuations in areas free of radio emission. We compared these rms errors (see Table 2) with their ideal values (0.032, 0.028 and 0.023 mJy in L, C and X band respectively) from the VLA manual (Perley, 2000). While in C and X bands the rms errors were found similar to their ideal limits, in L band they were larger than their theoretical values. This can be explained by additional interference at 20 cm, and is not a particularly unexpected result.

For 3 sources (MS0122.1+0903, MS1019.0+5139, MS2-347+1924) at 20 cm we measured a total flux density smaller than the peak flux density (but consistent within

the errors) probably because of phase errors not completely removed from radio maps by cleaning processes. Due to the small numerical difference between F_{tot} and F_{peak} (F_{tot}/F_{peak} equal to 0.93, 0.95 and 0.92 respectively) and since BL Lacs are known as objects with a high core-to-extended ratio (e.g. Perlman & Stocke, 1993), this does not introduce significant biases.

In three sources (namely MS0607.9+7108, MS1443.5+6348, MS1552+2020) we measured a very high rms threshold because of a source in the neighborhood, either very bright or extended, that could not be properly cleaned. In Table 3 we report simultaneous data for the 4 objects included in the CWSR sample from Stocke et al. (1985). The total flux density F_{tot} and their associated rms errors are listed at all the observed wavelengths.

3.2.2. Non-simultaneous flux densities

Name	NVSS * (20cm)		EMSS * (6cm)
	F_{tot}	F_{peak}	F_{tot}
MS0122.1+0903	2.49	1.75	1.4
MS0158.5+0019	12.7	12.7	11.3
MS0205.7+3509	4.7	4.7	3.6
MS0257.9+3429	10.1	10.1	10.0
MS0317.0+1834	20.9	17.9	11.5
MS0419.3+1943	2.85	2.76	8.0
MS0607.9+7108	25.2	25.2	18.2
MS0737.9+7441	22.7	22.7	24.0
MS0922.9+7459	84.2	59.2	3.3
MS0950.9+4929	2.6	2.6	3.3
MS1019.0+5139	4.8	4.8	2.4
MS1050.7+4946	64.1	64.1	53.8
MS1207.9+3945	18.7	18.7	6.1
MS1221.8+2452	24.5	24.5	26.4
MS1229.2+6430	58.0	54.5	42.0
MS1235.4+6315	12.1	12.1	7.0
MS1402.3+0416	40.7	40.7	20.8
MS1407.9+5954	35.7	33.3	16.5
MS1443.5+6348	17.6	17.6	11.6
MS1458.8+2249	31.9	31.9	29.8
MS1534.2+0148	72.3	72.3	34.0
MS1552.1+2020	79.0	79.0	37.5
MS1757.7+7034	10.3	10.3	7.2
MS2143.4+0704	101.1	95.1	50.0
MS2336.5+0517	16.5	13.6	4.9
MS2347.4+1924	4.7	4.7	3.2

* Total and Peak Flux Density in mJy

Table 4. Total flux densities drawn from NVSS (20 cm) and EMSS (6 cm) archives for the 26 sources. NVSS peak flux density is also listed.

Most datasets available to compute radio spectra are not simultaneous; moreover, BL Lacs are known to be variable objects (e.g. Angel & Stockman, 1980; Ulrich, Maraschi & Urry, 1995). To assess the impact of variability on non-simultaneous measurements of α_r , we statistically compare the α_r values measured from simultaneous data to values determined with non-simultaneous data. We used 20cm NVSS (Condon et al. 1998) and 6cm EMSS (Maccacaro et al. 1994) catalog fluxes. Since MS0122.1+0903 and MS0419.3+1943 do not appear in the NVSS catalog (they are at the completeness limit) we measured their fluxes from the NVSS on-line radio

Name	L band			C band			X band		
	F_{tot}^a	F_{peak}^a	rms($\times 10^{-2}$) ^b	F_{tot}^a	F_{peak}^a	rms($\times 10^{-2}$) ^b	F_{tot}^a	F_{peak}^a	rms($\times 10^{-2}$) ^b
MS0122.1+0903	1.34	1.44	6.26	1.47	1.39	3.30	1.28	1.08	2.60
MS0158.5+0019	13.26	12.46	8.67	9.70	8.87	3.00	8.16	7.36	3.00
MS0205.7+3509	3.53	3.53	13.85	5.24	5.04	3.40	4.96	4.24	2.49
MS0257.9+3429	8.80	8.23	10.00	10.79	9.41	3.80	11.23	9.98	3.05
MS0419.3+1943	6.10	5.66	8.44	9.05	7.91	3.18	8.07	7.16	2.57
MS0607.9+7108	15.33	13.12	40.71	14.37	13.97	3.58	10.34	9.82	3.74
MS0737.9+7441	25.60	24.78	9.29	22.46	21.04	3.45	20.07	19.39	3.26
MS0922.9+7459	43.37	21.88	8.57	4.17	3.59	4.00	2.89	2.66	3.30
MS0950.9+4929	4.58	4.56	5.67	3.42	3.16	2.88	3.25	2.89	2.72
MS1019.0+5139	2.71	2.84	8.03	2.97	2.75	3.16	2.86	2.70	2.60
MS1050.7+4946	57.13	48.84	7.34	42.71	39.59	3.46	38.19	32.44	3.31
MS1221.8+2452	23.31	21.17	13.02	22.00	20.95	3.50	21.59	19.80	2.75
MS1229.2+6430	54.68	51.51	11.87	46.31	43.97	4.65	40.96	38.21	3.20
MS1407.9+5954	31.51	24.54	12.67	19.55	16.80	4.44	16.77	13.67	3.02
MS1443.5+6348	9.74	9.55	51.29	4.47	3.88	4.69	6.29	5.11	3.17
MS1458.8+2249	48.42	47.47	8.75	87.99	85.33	5.23	94.41	91.87	3.00
MS1534.2+0148	65.12	60.16	8.10	34.07	25.29	3.29	24.01	15.96	2.57
MS1552.1+2020	64.46	56.92	53.89	38.79	31.93	3.79	29.61	24.06	2.71
MS1757.7+7034	11.68	11.45	15.61	11.15	10.53	3.78	9.80	9.03	3.07
MS2143.4+0704	78.59	65.10	15.48	51.87	45.27	6.31	43.57	37.21	2.29
MS2336.5+0517	11.82	8.83	12.62	5.86	5.26	3.30	4.69	4.38	2.80
MS2347.4+1924	4.27	4.64	9.96	3.01	2.87	2.90	2.42	2.19	2.70

^a Total and Peak Flux Density in mJy

^b Root Mean Square in mJy \times beam⁻¹

Table 2. F_{tot} , F_{peak} and rms measured in the L (20 cm), C (6 cm) e X (3.6 cm) frequency band for the 22 XBLs of CWSR sample observed on June 2000.

Name	F_{tot} (mJy)				
	20cm	18cm	6cm	2cm	1.3cm
MS0317.0+1834	17.3±0.6	16.5±0.8	11.5±0.23	5.5±0.3	<4.5
MS1207.9+3945	15.1±1.1	18.3±1.3	6.1±0.4	0.9±0.3	<4.2
MS1235.4+6315	10.7±0.5	10.1±0.5	7.0±0.38	4.6±0.3	Not Obs.
MS1402.3+0416	35.5±1.3	38.2±1.4	20.8±0.4	14.5±0.3	8.7±1.4

Table 3. F_{tot} and rms reported from Stocke et al. (1985) for the 4 EMSS objects included in the CWSR sample.

images which can be down-loaded in FITS format (see <http://www.cv.nrao.edu/NVSS/postage.html>).

In Table 4 flux densities are shown for the 26 listed objects.

3.3. Radio Spectra

By using the multi-frequency flux densities measured (Tables 2 and 3) we construct simultaneous spectra for CWSR sources. The spectral shape is assumed to be a power law ($S_\nu \propto \nu^{-\alpha_r}$). The slopes were computed with the Least Square Method (LSQ) and are reported in Table 5. At 20 cm the MS0922.9+7459 radio map is confused due to the superposition of another source (probably a radio galaxy in a foreground cluster). Thus, the corresponding flux density reported in this work, even if it is the best measure we could obtain, is to be considered as an upper limit and was not included in all the spectral slope calculations for this source.

To investigate whether variability could affect the spectral slope, we calculated simultaneous and non-simultaneous two-point (20 cm + 6 cm) α_r values with the EMSS (6 cm), NVSS (20 cm) and June 2000 (20 cm + 6 cm) observations for the CWSR sources.

We computed as well the spectral index for the source cores, finding no significant changes with respect the α_r calculated with total flux density. The only source that presents a remarkable difference between the core spectral

Name Name	sim. α_r	non-sim. α_r		
	CWSR -	NVSS/ EMSS	NVSS/ CWSR	CWSR/ EMSS
MS0122.1+0903	0.09	0.48	0.44	0.02
MS0158.5+0019	0.28	0.10	0.22	0.13
MS0205.7+3509	-0.22	0.22	-0.09	-0.017
MS0257.9+3429	-0.15	0.01	0.05	-0.11
MS0317.0+1834	0.48	0.57	0.50	0.02
MS0419.3+1943	-0.19	-0.86	-1.00	-0.23
MS0607.9+7108	0.20	0.27	0.47	-0.14
MS0737.9+7441	0.14	-0.05	0.01	0.05
MS0922.9+7459	0.72	2.70	2.50	2.32
MS0950.9+4929	0.21	-0.20	-0.23	0.27
MS1019.0+5139	-0.01	0.56	0.40	0.14
MS1050.7+4946	0.24	0.15	0.34	0.05
MS1207.9+3945	1.25	0.94	0.93	0.80
MS1221.8+2452	0.05	-0.06	0.09	-0.10
MS1229.2+6430	0.16	0.27	0.19	0.22
MS1235.4+6315	0.36	0.48	0.46	0.35
MS1402.3+0416	0.42	0.59	0.56	0.64
MS1407.9+5954	0.37	0.64	0.50	0.54
MS1443.5+6348	0.33	0.35	1.14	-0.15
MS1458.8+2249	-0.41	0.06	-0.84	0.40
MS1534.2+0148	0.57	0.63	0.63	0.54
MS1552.1+2020	0.45	0.62	0.59	0.45
MS1757.7+7034	0.09	0.30	-0.07	0.40
MS2143.4+0704	0.34	0.59	0.55	0.38
MS2336.5+0517	0.53	1.01	0.86	0.73
MS2347.4+1924	0.39	0.32	0.37	0.31

Table 5. Simultaneous and non-simultaneous spectral indices are listed for 26 objects ($S_\nu \propto \nu^{-\alpha_r}$).

index ($\alpha_r = 0.14$) and the total spectral index ($\alpha_r \leq 0.72$) is MS0922.9+7459.

In Table 5, spectral indices are listed from non-simultaneous and simultaneous spectra. The second col-

umn gives simultaneous slopes for the 26 sources of the CWSR sample; in the third, fourth and fifth column non-simultaneous slopes are computed, respectively, with 20cm NVSS and 6cm EMSS observations, with 20cm NVSS and 6cm CWSR observations, with 20cm CWSR and 6cm EMSS observations. By comparing different values of spectral slope, we note that 11 of 26 objects have spectral indices that could be classified differently depending on the dataset used; in particular, 9 of 11 sources could have a flat or steep spectrum, 1 source (MS0419+19-43) could have an inverted or flat spectrum, 1 source (MS1458.8+2249) could have an inverted, flat or steep spectrum.

In the next section we will consider a few explanation that could affect these measurements: the well-known variability of these objects and the angular resolution of the observations. Data taken with different angular resolution in fact could measure, if present, a different amount of extended flux.

4. Results

4.1. Core-to-total flux density ratio

frequency band	α	β
L (20 cm)	0.94	0.02
L (20 cm) *	0.96	0.01
C (6 cm)	0.99	-0.03
X (3.6 cm)	1.00	-0.06

*Fit without MS0922.9+7459.

Table 6. Results of the fit of F_{Peak} vs. F_{Tot} for the CWSR sources indicating that an extended component is visible at 20 cm.

It is already known that BL Lacs are very core-dominated sources (e.g. Perlman & Stocke, 1993) and that this behavior is more evident in higher frequency observations. This is confirmed from our data (see section 3.2.1). In fact, we confirm the trend of the core-dominance in the CWSR sample by fitting F_{Peak} vs F_{tot} (from Table 2) with the LSQ. The LSQ slope and intercept are given in Table 6: the first column gives the VLA frequency band, the second and the third column, respectively, give the slope and the intercept of the LSQ fit. We compute the fit also excluding the L band value of MS0922.9+7459 (that is confused, see section 3.3)

The errors on the slopes are typically 0.1-0.2, and therefore only at 20 cm there is a small deviation from a linear fit, indicating the presence of an extended component which is more visible for brighter sources.

We also checked for a possible redshift dependence of the core-dominance since extended surface brightness dims as $(1+z)^{-4}$, so significant amount of flux from the extended component could be lost at high z . In addition, since at large distances a given angular resolution corresponds to a larger physical size, if much of the extended

flux of a source is close to the core, some of its inner part could be subsumed into the central component and therefore included with the measured core flux. Perlman & Stocke (1993) studied the possible biases introduced by these terms and found that for sources with $z \leq 0.5$ there is only a modest correction. Since the objects in this work are all at $z \lesssim 0.5$, we estimate that the loss of the extended flux due to these biases was not significant and, thus, was not considered.

4.2. Spectral indices

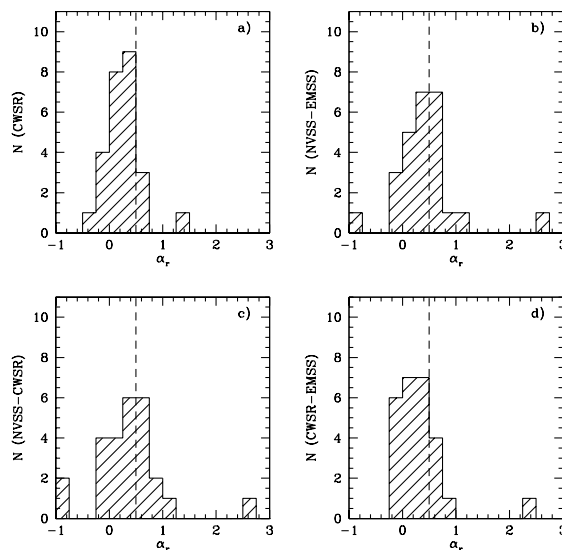


Figure 1. The α_r distributions for the 26 sources in the CWSR sample. Plot a), upper left, shows the spectral indices calculated from the simultaneous CWSR data. Plot b), upper right, shows the spectral indices calculated from 20cm NVSS data and 6cm EMSS data. Plot c), bottom left, shows the spectral indices calculated from 20cm NVSS data and 6cm CWSR data. Plot d), bottom right, shows the spectral indices calculated from 20cm CWSR data and 6cm EMSS data. The vertical dashed line represents the cut-off value $\alpha_r = 0.5$.

We studied the distribution of spectral slopes reported in Table 5. In Figure 1 the 4 α_r distributions obtained for the simultaneous and non-simultaneous datasets are shown. The dividing line at $\alpha_r = 0.5$ is plotted as a dashed line.

In Table 7 we report the percentage of sources with a spectrum steeper than $\alpha_r = 0.5$. The first column gives the different datasets used to compute the spectra; the second column gives the percentage of steep spectra computed; the third column gives the time span between the different datasets; the fourth column gives the VLA configurations used for the observations. As shown in Table 7, we found that about 15% of the sources have steep spectra com-

datasets		%	Δ	VLA
20 cm	6 cm	steep sp.	yrs	config
CWSR	CWSR	15.4	0	C+C/D*
NVSS	EMSS	38.5	5-10	D, C
NVSS	CWSR	38.5	4-7	D, C
CWSR	EMSS	23.1	12-14	C+C/D, C

* Also 3.6 cm observations in C configuration and 18 and 2 cm observations in C/D configuration are used

Table 7. The percentage of objects with steep spectra ($\alpha_r > 0.5$; $S_\nu \propto \nu^{-\alpha_r}$) is listed for different datasets.

puted with simultaneous data, and even larger percentages with non-simultaneous data. As shown in Figure 1 the distribution of α_r is a rather continuous one, so that picking a single value to divide it in two populations is rather arbitrary, and does not have a physical meaning. However, this has both practical and historical reasons (see e.g. the 1Jy BL Lac sample definition in Kühr & Schmidt, 1990) and so it remains an important issue to be addressed.

To investigate whether the different subsamples under consideration are consistent with each other, we performed a Kolmogorov-Smirnov test for the four α_r distributions (see Figure 1) by comparing their respective datasets considered in Table 7. We cannot exclude that the four dataset combinations are extracted from the same population ($P > 32\%$). However, the small number of points does not allow any further investigations.

Since by using non-simultaneous data we find a higher percentage of steep spectrum sources, it seems straightforward to ascribe this effect to flux variability, that we know is present in this class of objects. However, other factors are at play. Again from Table 7, we notice that the higher percentages of steep spectra correspond to the use of the 20cm NVSS data, which were carried out with the VLA array in D configuration. As previously discussed (see 4.1), this probably indicates that a larger fraction of extended flux, which has a spectrum steeper than the core emission, is detected in these datasets. The Gaussian distribution of the NVSS-to-CWSR 20cm data ratio seems to confirm this hypothesis: the distribution mean value is $x_{mean} = 1.19$, indicating that the NVSS flux densities exceed on average the CWSR values by 19%.

The effect cannot be ascribed to confusion due to a large beam, since inspection of the higher resolution C configuration images shows that no sources fall into the NVSS beam beside the target BL Lac.

Thus, along with variability, we might suggest that a significant fraction of extended flux, that is considered negligible by BL Lacs beaming models (e.g. Blandford & Rees, 1978, Antonucci & Ulvestad, 1985), is present, and could at least partially explain the differences obtained in the α_r distributions.

To investigate what influence the detection of an extended component has on the spectral index, we plotted the spectral slopes of the CWSR sources against redshift z in Figure 2. We confirm what we assumed in §4.1 that no redshift effect is present in the detection of extended flux due to decreasing of brightness. The LSQ fit in Figure 2 is consistent with a constant due to the large error on

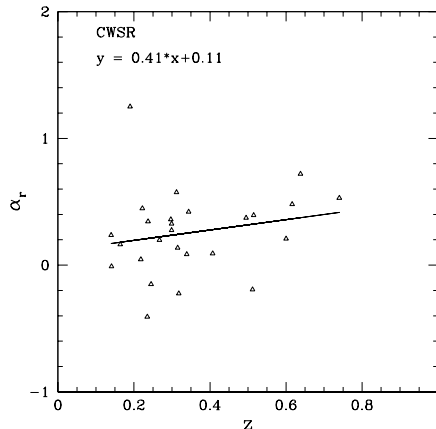


Figure 2. α_r against z is plotted for the simultaneous dataset. The LSQ fit and the respective functional form are also shown.

slope (0.41 ± 0.39) and if anything sources at larger redshift seem to have steeper slopes.

To assess possible biases due to the varying spectral slope, we computed total luminosity for the 25 XBLs with known redshift in the L, C and X frequency bands. For the 22 CWSR sources observed on June 2000, peak and extended luminosities could also be computed from the peak and total flux densities; where extended emission was not detected, conservative upper limits on extended radio luminosity levels were obtained by assuming that each source has uniformly bright extended emission at 1σ detection level over a 3000 kpc^2 area surrounding the core (see Rector et al. 2000 for a description of how to compute the extended luminosities). The luminosity of the extended component is in general in agreement with what found previously (e.g. Rector et al. 2000, Perlman & Stocke 1993) so we do not report values here. In fact the C configuration snapshot data do not have a sufficient uv coverage to warrant a better definition of the core than the previous A configuration data on the same sources.

To test whether there is a correlation between luminosities and spectral slope, we computed the correlation coefficient for the LSQ fit and its goodness probability. We found again that there is a 99% probability that a correlation exists only for extended luminosity at 20cm. We check also a possible dependence of α_r on NVSS and EMSS radio luminosities. The correlation between spectral slope and 20cm luminosity seems to support the idea that the extended component is important in these objects

and that this is the main factor to explain the difference in the spectral measures.

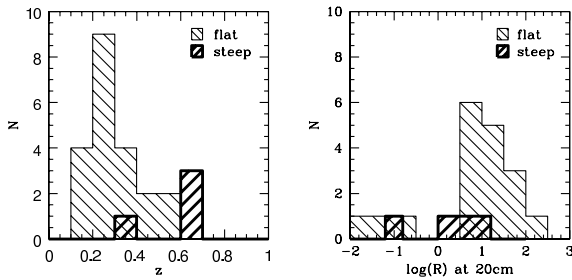


Figure 3. The redshift *left* and core dominance *right* distribution of flat spectra objects compared to the steep spectra ones.

By splitting our sample of 26 objects in two groups, steep and flat spectra, we studied the different distribution of redshift, radio luminosity (at 20 and 6 cm), X luminosity (0.3-3.5 keV) and 20cm core dominance.

We plot in Figure 3(*left*) the histogram of the redshift distribution for objects with $\alpha_r > 0.5$ and $\alpha_r \leq 0.5$ respectively. The flat-spectrum dataset has on average redshift values smaller than the steep-spectrum one. Since this is a flux limited sample, it is fair that we find the same behavior also for the radio and X-ray luminosities. However we stress that the redshift range of this sample is small, and in particular lower than the radio selected samples like the 1Jy.

In Figure 3(*right*), we show the histogram of core dominance for the two datasets, at 20 cm. As expected, the steep spectrum sources show in general a smaller degree of core dominance, even if the statistical significance of the result is marginal.

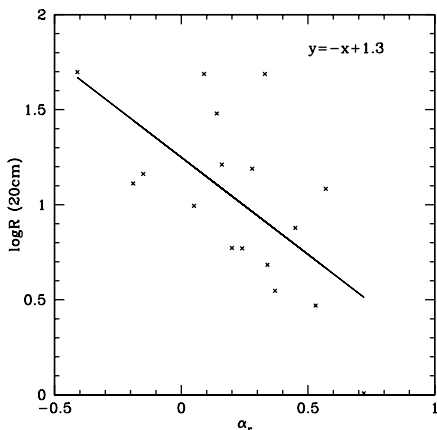


Figure 4. The logarithm of core-dominance R (at 20 cm) is plotted against the radio spectrum α_r . The regression line has a slope of -1.02 with rather large uncertainties (0.32).

We also plot in Figure 4 the core-dominance at 20 cm against the radio spectral slope α_r , for the sources for which we have a measure (i.e. not an upper limit) for R. The two variables have a correlation coefficient of -0.63 , indicating that, within the available statistics, they are correlated at $> 99\%$.

4.3. $\langle V_e/V_a \rangle$ and the spectral index selection

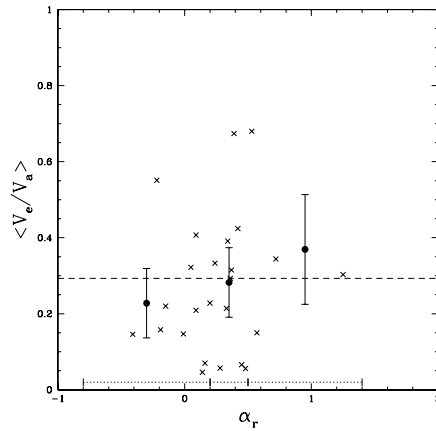


Figure 5. V_e/V_a against α_r is shown for the 25 XBLs with a redshift measure. The $\langle V_e/V_a \rangle$ trend is also plotted with error-bars for α_r grouped in three bins ($\alpha_r < 0.2$; $0.2 < \alpha_r < 0.5$; $\alpha_r > 0.5$). The dotted line represents the mean value of $\langle V_e/V_a \rangle$.

One controversial issue in the study of BL Lacs is their evolutionary behavior. The first studies were based on the EMSS and found evidence for “negative evolution” (i.e. less numerous/luminous objects in the past; Maccacaro et al. 1984; Morris et al. 1991; Wolter et al. 1994; Rector et al. 2000). At the same time the 1Jy sample (selected in the radio band, and with a constraint on $\alpha_r < 0.5$, Kühr & Schmidt 1990) showed no or mild positive evolution (Stickel et al. 1991; Rector & Stocke 2001). It is therefore interesting to think that the different criteria, in particular the flat radio spectral index, could select samples with different evolution properties. This result is obtained if the steep spectrum BL Lacs, present only in X-ray selected samples, all show a very negative evolution. Since the number involved in these samples are small (20-40 objects in total), even few objects have a large weight.

We used the standard $\langle V_e/V_a \rangle$ test (Avni & Bachall, 1980) to check the uniformity of the spatial distribution of the objects in a flux limited sample. For each object the two volumes (V_e and V_a) are derived on the basis of the observed redshift and the maximum redshift at which the object could be still included in the survey. A value of $\langle V_e/V_a \rangle = 0.5$ is given by a uniform distribution, while lower values of $\langle V_e/V_a \rangle$ indicate a distribution skewed towards low redshifts. Details can be found e.g. in Morris et al (1991). We stress here that only the X-ray flux limits

are used to derive the maximum available volume. In fact the EMSS did not have any requirement on the radio flux or optical flux to have an object enter the sample, and all the EMSS BL Lacs have been also detected in the radio band.

To test the impact of the α_r selection criteria on the measure of the evolution, we studied the correlation between V_e/V_a and the radio spectral index α_r for the 25 objects in the CWSR sample with a redshift measure (see Rector et al. 2000 for the V_e/V_a values), plotted in Figure 5. The data have been also grouped in three α_r bins to better show the trend ($\alpha_r < 0.2$; $0.2 < \alpha_r < 0.5$; $\alpha_r > 0.5$) and plotted in Figure 5 with error bars that reflect number statistics. We find that the mean value is consistent with a constant in the 3 intervals of α_r . If anything there is a slight suggestion that XBLs with steep spectra ($\alpha_r > 0.5$) have a higher mean $\langle V_e/V_a \rangle$ value. Thus, it seems that the exclusion of steep spectra objects in RBLs samples could not cause the differences observed in the cosmological evolution with respect to the EMSS sample.

The picture is however far from clear: while surveys like the Deep X-ray Radio Blazar Survey (DXRBS; Perlman et al. 1998; Padovani et al. 2002; Giommi et al., 2002), the “sedentary” survey (Giommi et al. 1999) and the ROSAT - All Sky Survey (RASS; Bade et al. 1998; Beckmann et al. 2003) found very low values of $\langle V_e/V_a \rangle$ for XBLs, other surveys like the Radio Emitting X-ray survey (REX; Caccianiga et al. 1999, Caccianiga et al., 2002) did not find the same result, possibly indicating that the negative evolution in this class of objects is not so strong as was previously measured.

5. Conclusions

By using fluxes at 20, 6 and 3.6 cm, simultaneously measured with the VLA, we computed the radio spectra for a complete subsample of 22 XBLs from the EMSS (Gioia et al., 1990, Stocke et al., 1991). To this sample we added 4 sources whose simultaneous spectra were obtained by Stocke et al. (1985). The aim was to study the spectral slope distribution at radio frequencies without possible biases introduced by the flat-radio-spectrum criterion often used in selecting BL Lac samples. We found that about 15% of the sources have steep spectra. We considered the archived non-simultaneous data as well, finding an even higher percentage of steep spectra. This effect could be ascribed to variability, but other factors seem more significant in determining the measured slope of spectra. The higher percentage of steep spectra objects ($\sim 38\%$), obtained by using the 20cm NVSS observations, implies the presence of a large fraction of extended flux that is preferentially detected by the VLA D configuration. This seems also supported by a possible correlation between spectral index and 20cm extended luminosities, computed either with simultaneous or non-simultaneous data.

We find that the possible biases introduced by the flat-radio-spectrum criterion in the RBLs samples cannot easily account for the discrepancies observed in the $\langle V_e/V_a \rangle$

values of RBLs and XBLs samples and therefore the issue of evolution of different BL Lac classes is still open. The percentages of steep spectrum BL Lacs found for the EMSS sample cannot be applied straightforwardly to RBL samples, since a) the selection by radio spectral index is sometimes computed using non-simultaneous data; b) the XBLs properties are different from those of RBLs in the sense that RBLs are more variable and possibly more core dominated. If the observed spread in α_r is mostly due to variability, we should expect a larger fraction of steep RBLs than XBLs. If most of the effect is due, as the data seem to suggest, to the presence of a significant extended component, that is lesser in RBL, the influence on the radio selected objects should be less severe.

We stress here that the α_r distribution in the EMSS BL Lacs simultaneous data is a continuous one and so any division in two subsamples is somewhat arbitrary and does not have a physical underlying support. We are tempted to suggest a different cut-off value of $\alpha_r = 0.7$ (as done e.g. by Perlman et al. 1998 in DXRBS) that would then exclude $\leq 10\%$ of all BL Lacs based upon our findings here. But, because this would also include a very large number of “normal” radio galaxies, some of which with optical featureless spectra (see e.g. discussions in Rector & Stocke, 2001, Perlman et al. 1996), other information would have to be used to define them as BL Lacs (e.g. variable optical polarization).

However, we have shown that the selection effects applied by using a cut-off value in the radio spectral index of $\alpha_r = 0.5$ are not so severe. The steeper BL Lacs that are excluded from this criterion seem to have the same overall properties of the flat ones, especially for what concerns the cosmological distribution. We can therefore conclude that the discrepancies found by using different samples cannot be ascribed to the flat radio spectrum criterion.

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